

# Hubble Speed from First Principles

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## Distances in Cosmology

Distances are set by the metric that describes our space-time; in a cosmological context, this implies that their scaling with time is related to the expansion rate of the Universe.

Comoving distance:

$$\chi(z) = \frac{1}{H_0} \int_0^z \frac{d\tilde{z}}{E(\tilde{z})}$$

where I will call <u>cosmological model</u>:

$$E^{2}(z) \equiv \frac{H^{2}(z)}{H_{0}^{2}} = \Omega_{m}(1+z)^{3} + \Omega_{r}(1+z)^{4} + \Omega_{\text{DE}}X(z)$$

X represents all contributions to the expansion that are not matter or radiation, e.g. cosmological constant (X=I), DE, modified gravity, spatial curvature, etc..

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## Distances in Cosmology

Comoving distance is not a direct observable. There are two types of distances measured in an expanding Universe:

#### Luminosity distance

standardizable candles/sirens

$$m_B = M_B + 5 \log \left(\frac{d_L}{10 \text{pc}}\right) + 25$$
$$M_B = 5 \log H_0 - 5a_B - 25$$

$$m_B = 5 \log \left(\frac{H_0 d_L}{10 \text{pc}}\right) - 5 a_B$$

#### Angular diameter distance

BAO standard rulers (CMB peaks)





 $\theta_s$ 

$$\theta_s(z) = \frac{r_s}{d_A(z)}$$
$$\Delta z = H(z)r_s$$

#### **Distance Duality Relation**

In any metric theory of gravity, if photons propagate along null geodesics and obey number conservation:

$$\eta(z) \equiv \frac{d_L(z)}{(1+z)^2 d_A(z)} = 1$$

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Let us set  $\eta = 1$  and rearrange the DDR as follows:

$$H_0 = \frac{1}{(1+z)^2} \frac{H_0 d_L(z)}{H(z) d_A(z)} H(z)$$

H(z)d\_A(z): from a combination of line-of-sight and transverse BAO measurements, without the need of an external anchor for the sound horizon at drag epoch [7 data points from BOSS]

- Cosmic chronometers (CC) measurements provide H(z) data that is free from calibration and does not depend on the underlying cosmological model [a compilation of 30 measurements]
- **H\_Od\_L** can be obtained from measurements of SNIa without the need of calibration [Pantheon]

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$$\eta(z) \equiv \frac{d_L(z)}{(1+z)^2 d} \qquad H(z) = -\frac{1}{1+z} \frac{\Delta z}{\Delta t}$$
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Combining BAO and CC first, and then folding in SNIa, we can get a *calibration- and model-independent* estimate of the Hubble parameter:



$$H_0 = \frac{1}{(1+z)^2} \frac{\left[H_0 d_L(z)\right]^{\text{SNIa}}}{\left[d_A(z)\right]^{\text{BAO}+\text{CC}}}$$

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#### Some technicalities:

- **=** the three data set cover roughly same range,  $0 \leq z \leq 2.5$ , but data points are not exactly at same redshifts
- we choose the redshifts of BAO data points as reference ones, and perform GP fit to SNIa and CC data to get corresponding data points (including error bars), i.e. PDFs

- **a** at each of the 7 redshifts, we draw 10k realizations of  $[H_0 d_L]^{\text{SNIe}}$  and 10k realizations of  $[d_A]^{\text{BAO+CC}}$ . We combine them together following the DDR equation to get realizations (PDFs) of H<sub>0</sub> at each redshift.
- resulting mean values are consistent within error bars, but they are different at each redshift, with two being notably lower...(next slide!)
- we finally multiply the 7 PDFs to get one PDF for H from which we can extract mean and variance using the inverse transform sampling.
- Overall the method is a combination of GP and MCMC, we refer to it as GPMC.
- It is important to highlight that we work strictly within the redshift range of the data, without performing any extrapolation.



 $H_0 = 69.5 \pm 1.7 \,\mathrm{km/s/Mpc}$ 





correspondingly we infer the absolute magnitude to be  $M_B = -19.355 \pm 0.054$ i.e a shift of  $|\delta M_B| = 0.138 \pm 0.067$  wrt SH0ES value.

in agreement with  $\mu_{Cepheid} - \mu_{TRGB} = -0.139 \pm 0.024 \text{ mag}$  [Efstathiou, arXiv:2007.10716]

### A redshift trend in $H_0$ ?



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# A redshift trend in $H_0$ ?

Assuming the validity of the DDR, is equivalent to assuming that *the two probes are consistent tracers of cosmic expansion*.

If this was not the case, the H<sub>0</sub> inferred from the previous procedure would not be a constant, but it would rather show some, *unphysical trend in redshift*.



## perhaps a mild dynamical feature in H(z)



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